Neutrino oscillations in supernovae and mergers

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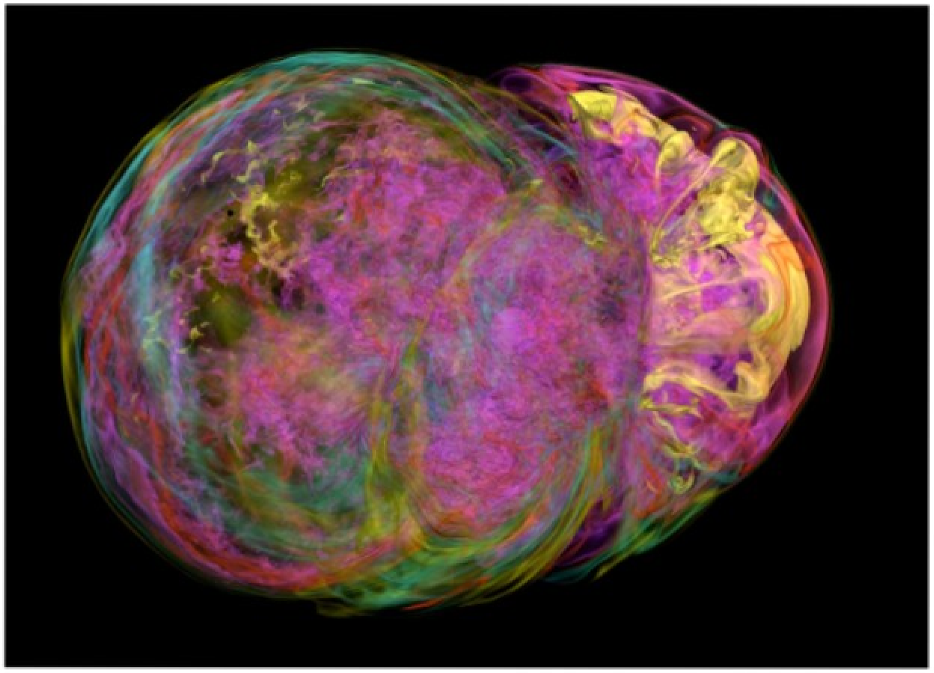
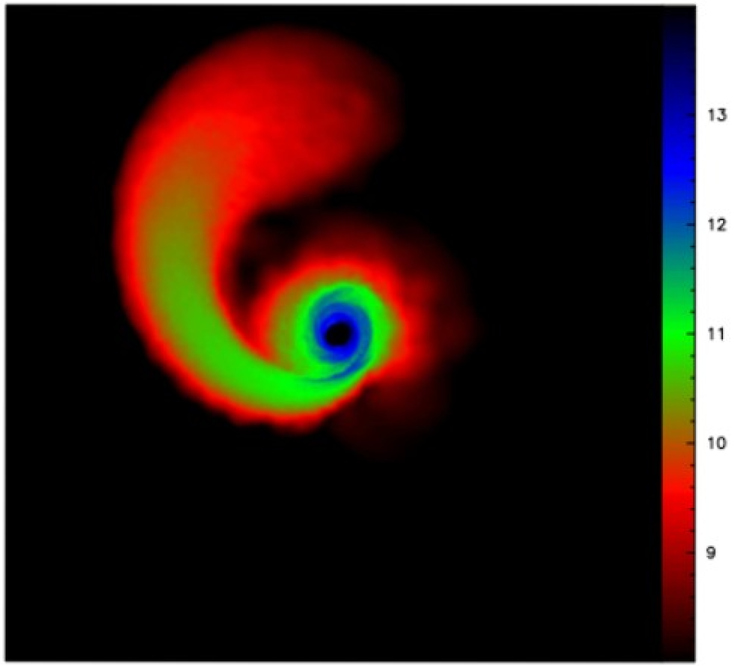
Core collapse supernovae and compact object mergers

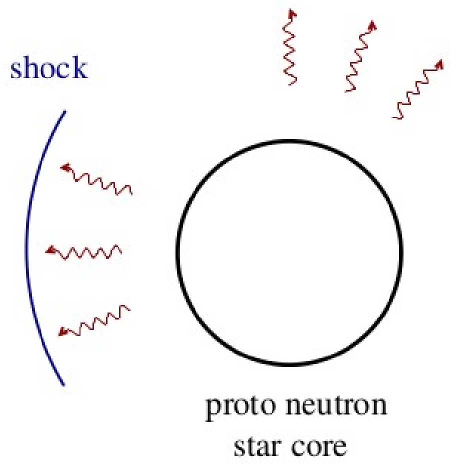


Core collapse supernova
(from Blondin et al.)

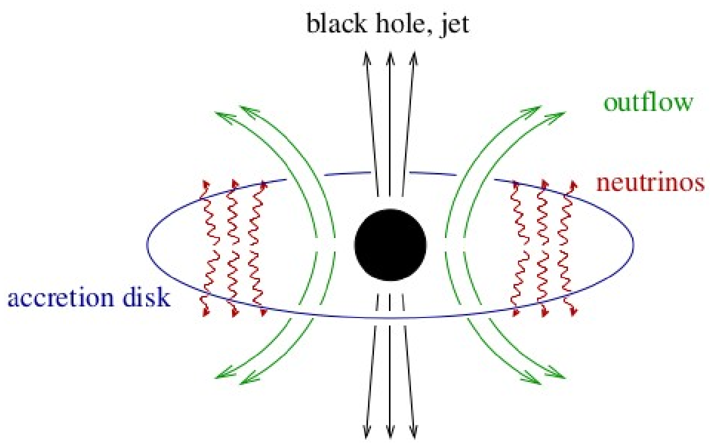


Compact object merger
(from Rosswog et al.)

Explosions and mergers

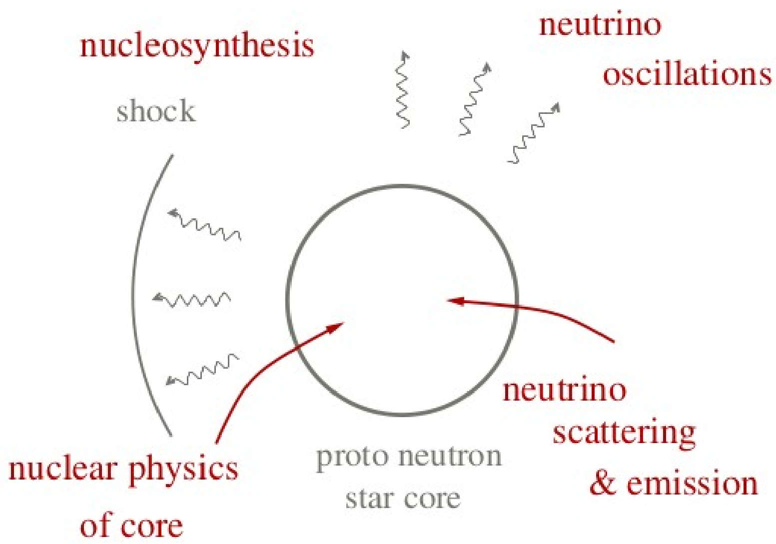


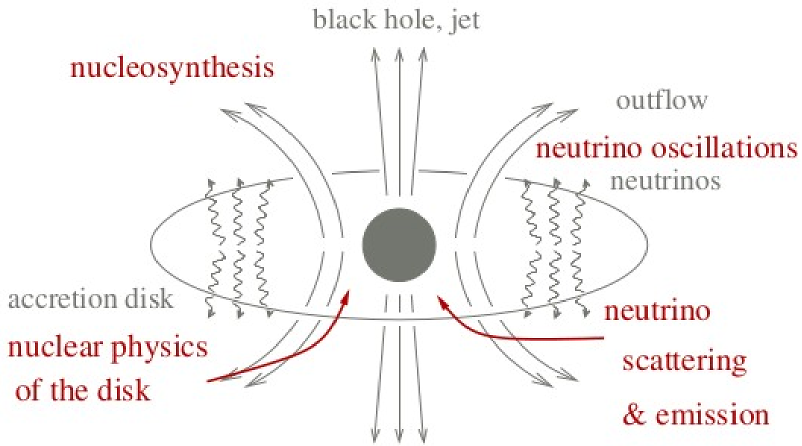
Standard core core collapse SN



Accretion disk from core collapse SN or compact object merger

Microphysics of explosions and mergers





Standard core core collapse SN

Accretion disk from core collapse SN or compact object merger

Specific examples of questions where neutrino physics is needed

How do neutrinos affect the dynamics of both objects?

What is the spectrum of a supernova neutrino signal?

Which r-process elements do neutron star mergers make?

What elements are made in supernovae winds?

The weak interaction matters

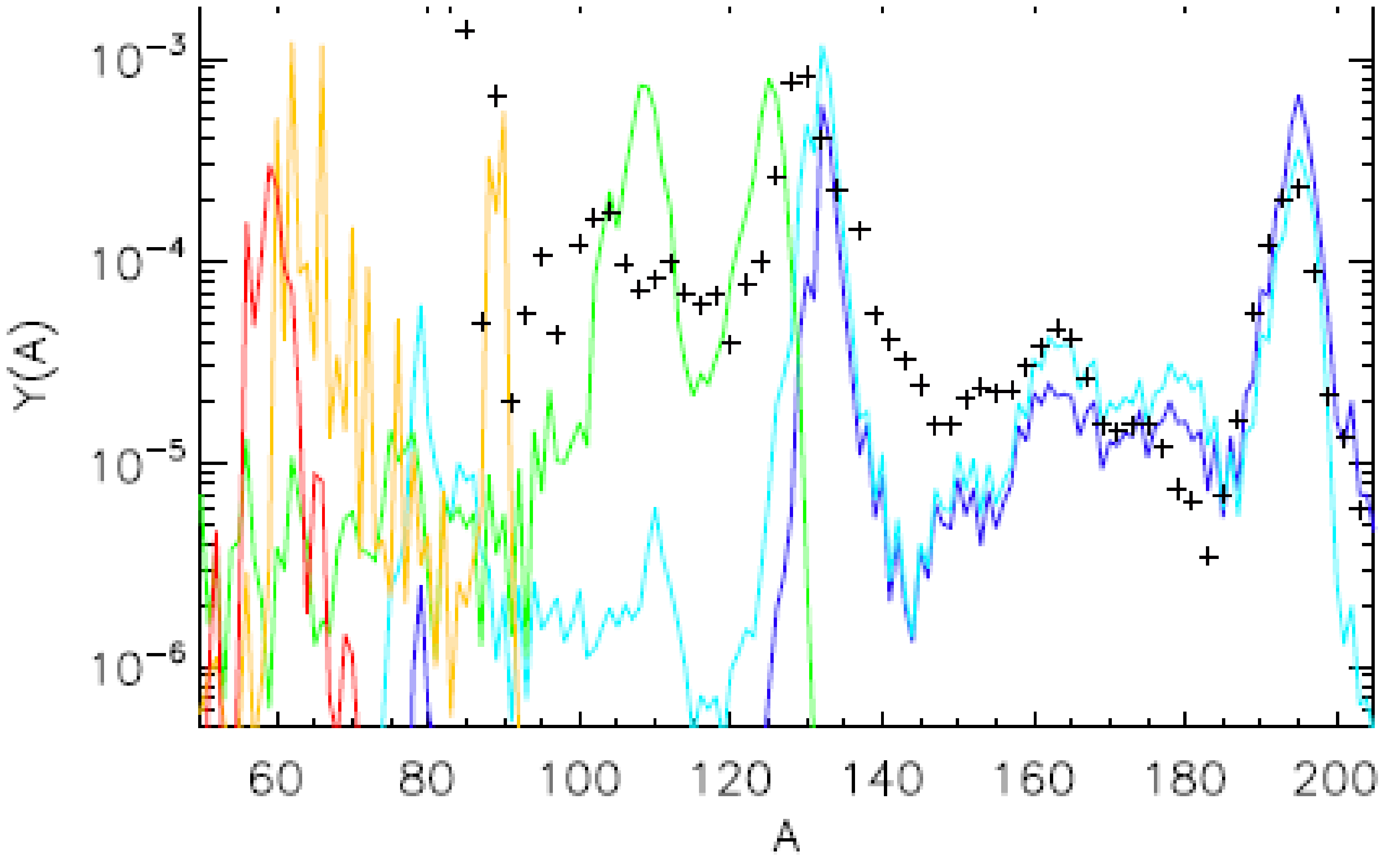
How neutrinos influence nucleosynthesis

Neutrinos change the ratio of neutrons to protons

$$\nu_e + n \rightarrow p + e^-$$

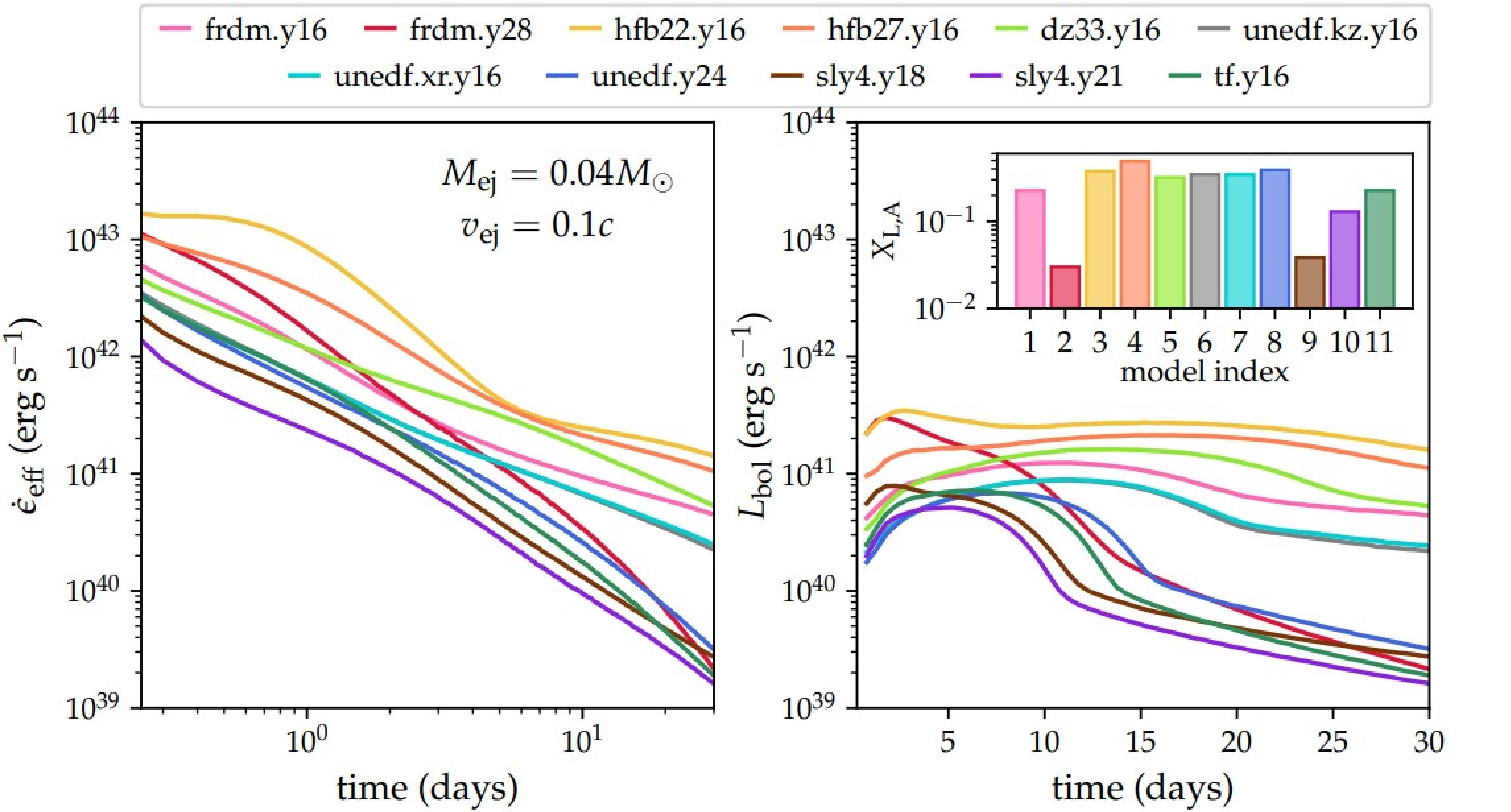
$$\bar{\nu}_e + p \rightarrow n + e^+$$

How much does it matter?



Malkus et al 2016

Electromagnetic counterpart (kilonova) to a merger depends on the electron fraction



Barnes et al 2021

Flavor matters for nucleosynthesis

Neutrinos change the ratio of neutrons to protons

$$\nu_e + n \rightarrow p + e^-$$

$$\bar{\nu}_e + p \rightarrow n + e^+$$

Oscillations change the spectra of ν_es and $\bar{\nu}_e$s

$$\nu_e \leftrightarrow \nu_\mu, \nu_\tau$$

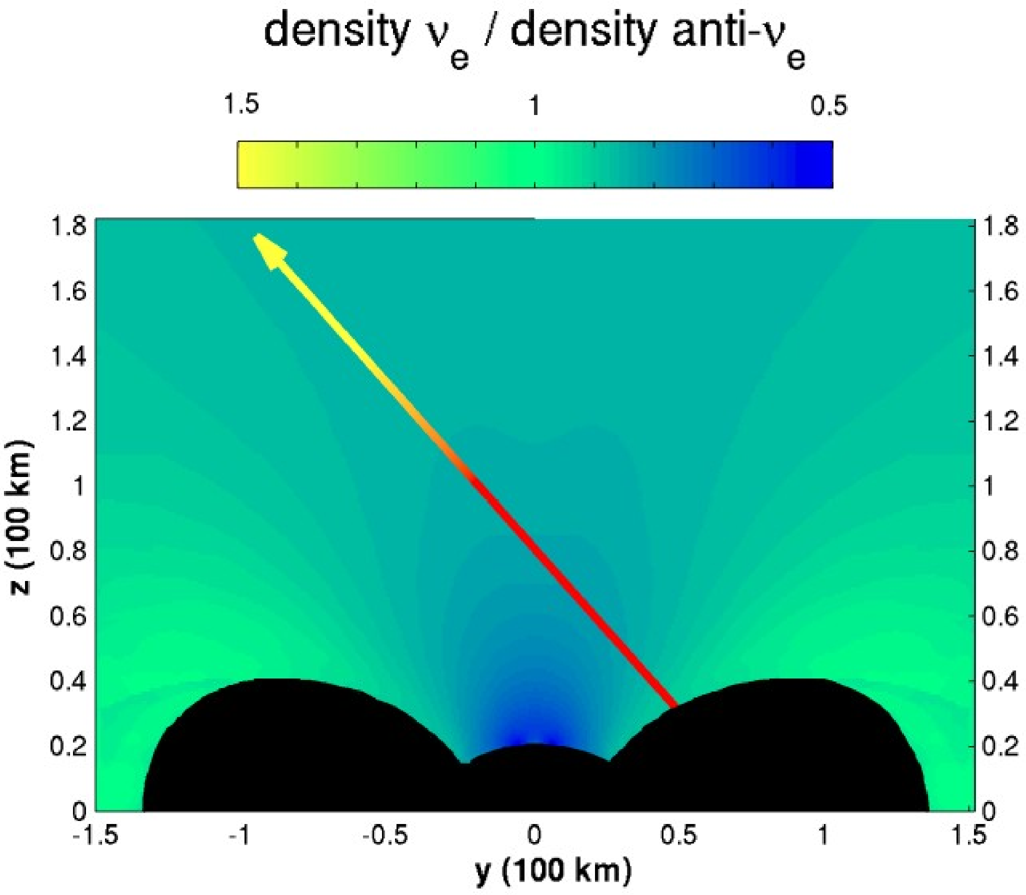
$$\bar{\nu}_e \leftrightarrow \bar{\nu}_\mu, \bar{\nu}_\tau$$

Mergers have less ν_μ, ν_τ than ν_e and $\bar{\nu}_e$

$\rightarrow$ oscillation reduces numbers of ν_e, $\bar{\nu}_e$

Do neutrinos transform in supernovae and mergers?

Answer, almost certainly, is yes



Zhu et al 2016

Neutrinos can be described by a density matrix

Additional information
about the phase

$$\rho = \begin{pmatrix} \rho_{ee} & \rho_{ex} \\ \rho_{ex}^{*} & \rho_{xx} \end{pmatrix}$$

Tells you how likely you are to
measure the neutrino as electron type

Tells you how likely you
are to measure the neutrino
In an x (mu or tau) state

Neutrinos can oscillate (flavor transform)

$$i\frac{D\rho}{Dt} = [\mathbf{H}, \rho] + iC$$

$$i\frac{D\bar{\rho}}{Dt} = [\bar{\mathbf{H}}, \bar{\rho}] + i\bar{C}$$

Collision term

Convective derivative

Hamiltonian

Hamiltonian creates non-linearity

$$\mathbf{H} = \mathbf{H}_{\rm vac} + \mathbf{H}_{\rm M} + \mathbf{H}_{\rm SI}$$
$$\bar{\mathbf{H}} = \mathbf{H}_{\rm vac} - \mathbf{H}_{\rm M} - \mathbf{H}_{\rm SI}^{*}$$

$$\imath \frac{D\rho}{Dt} = [\mathbf{H}, \rho]$$
$$\imath \frac{D\bar{\rho}}{Dt} = [\bar{\mathbf{H}}, \bar{\rho}]$$

Neutrinos see a potential due to other neutrinos

Neutrinos see a potential due to the matter

Flavor and mass are not the same

Types of transformations

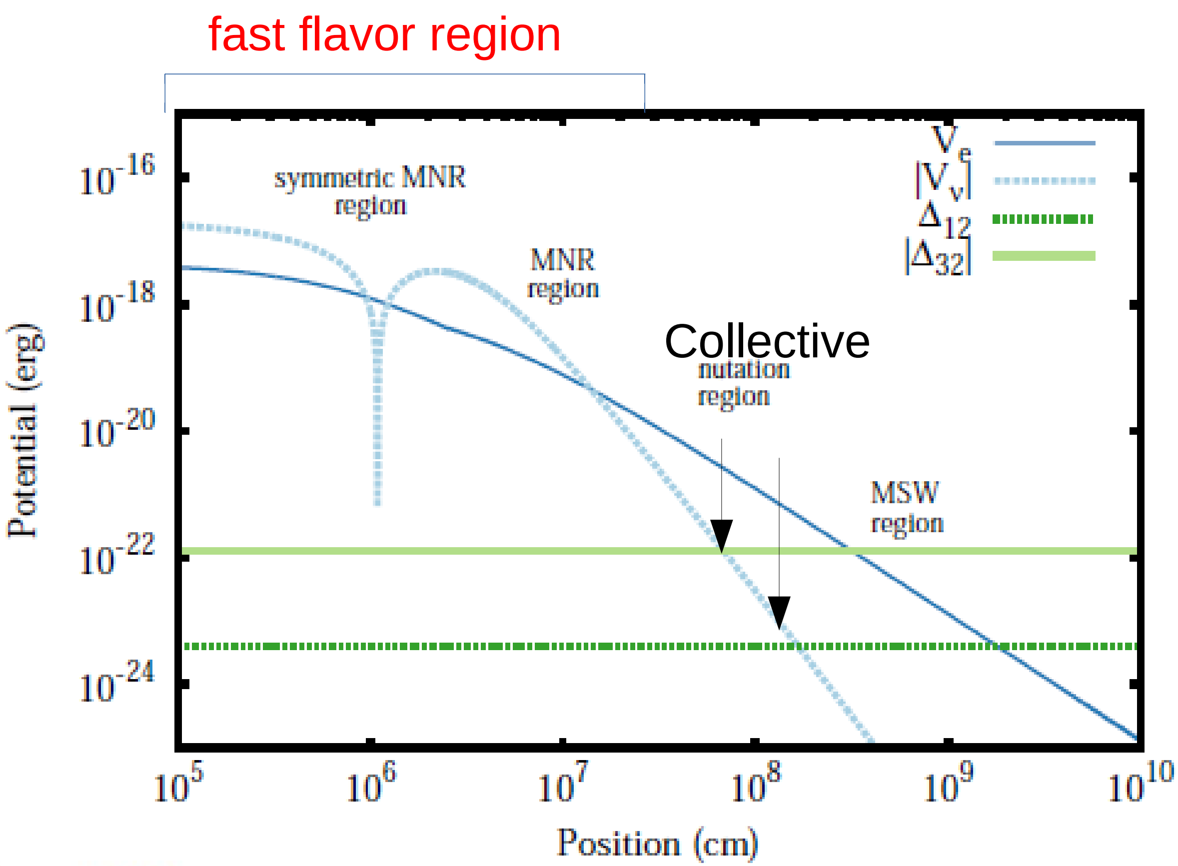
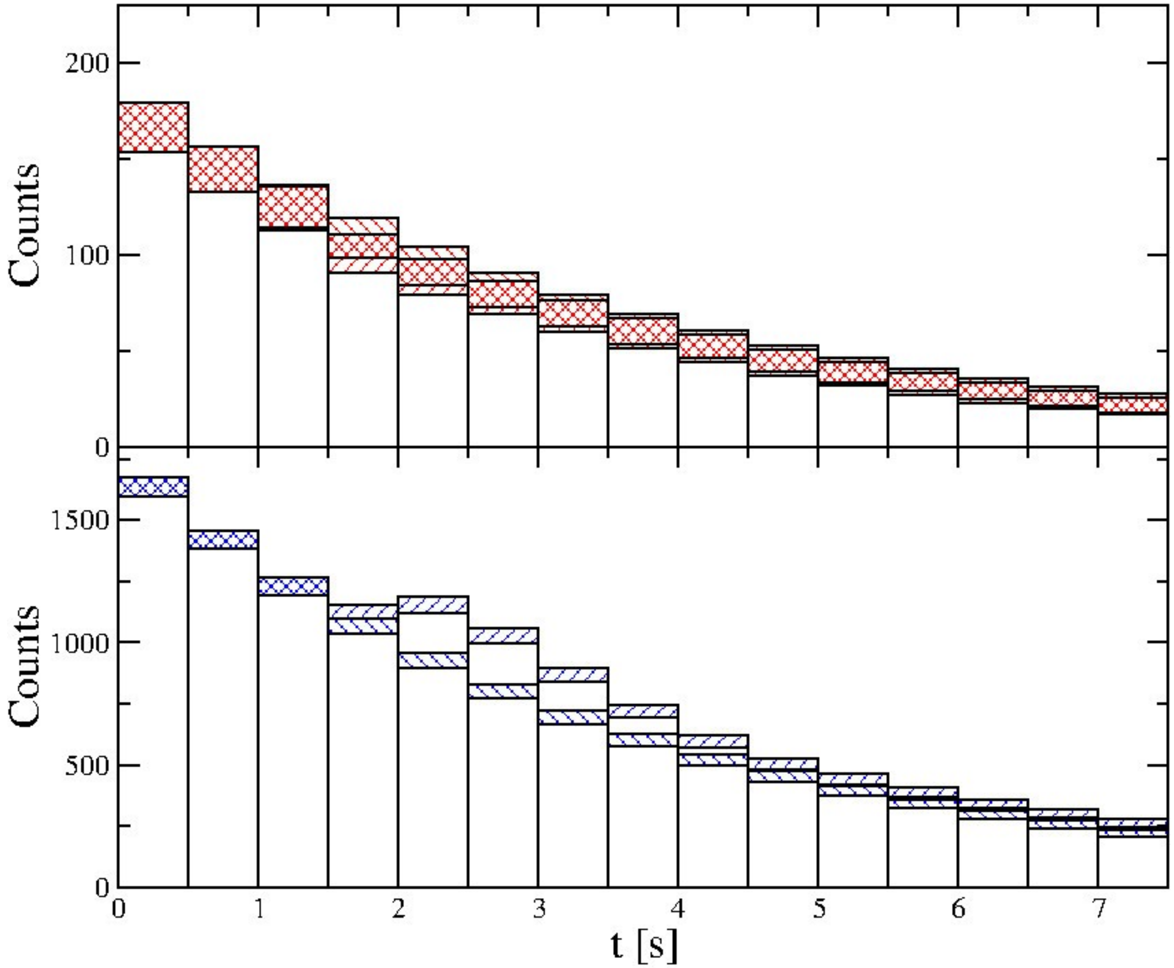


fig. from Malkus et al 2016

$$\mathbf{H} = \mathbf{H}_{\rm vac} + \mathbf{H}_{\rm M} + \mathbf{H}_{\rm SI}$$

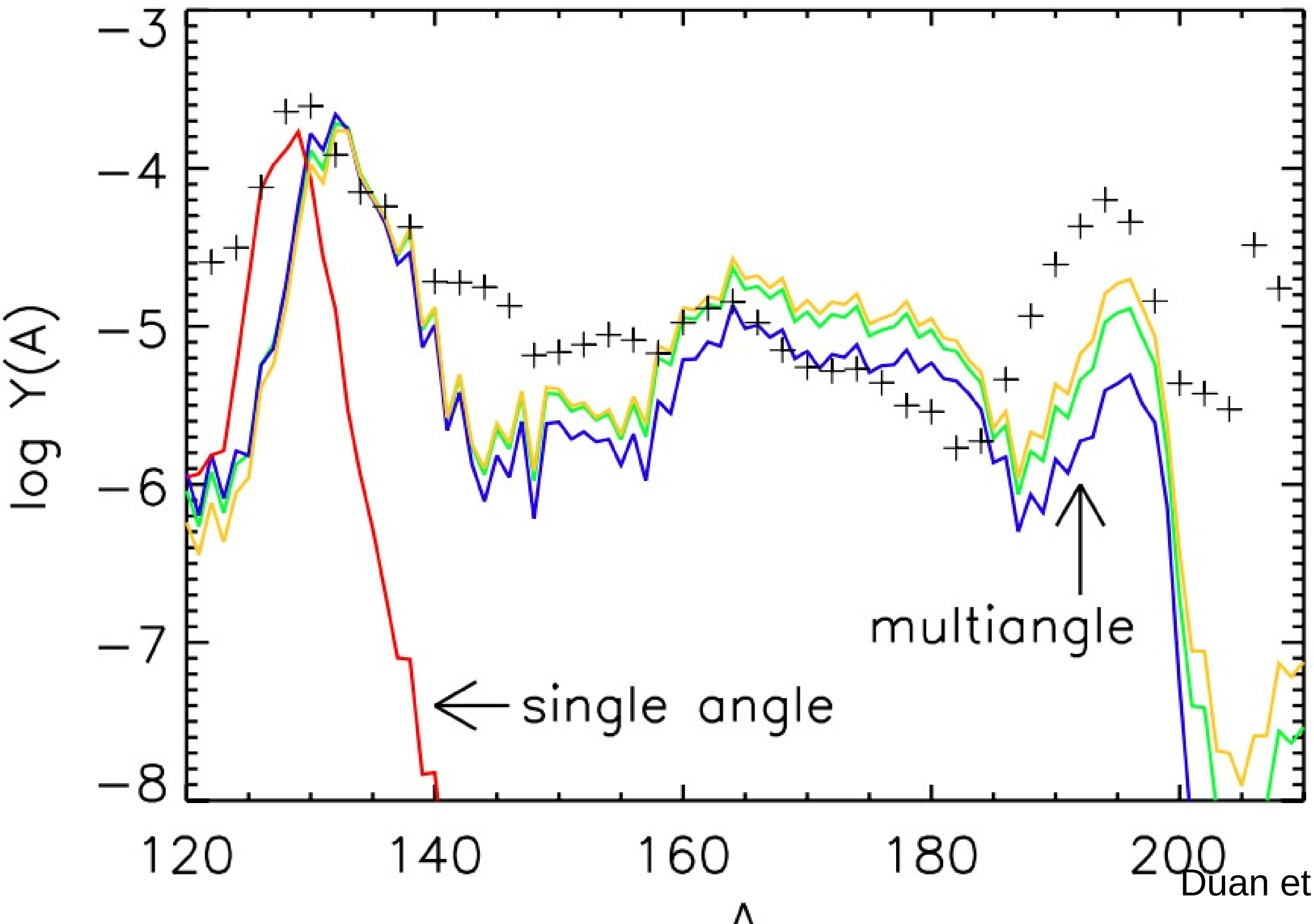
$$\bar{\mathbf{H}} = \mathbf{H}_{\rm vac} - \mathbf{H}_{\rm M} - \mathbf{H}_{\rm SI}^*$$

Collective and MSW change on the supernova neutrino time signal



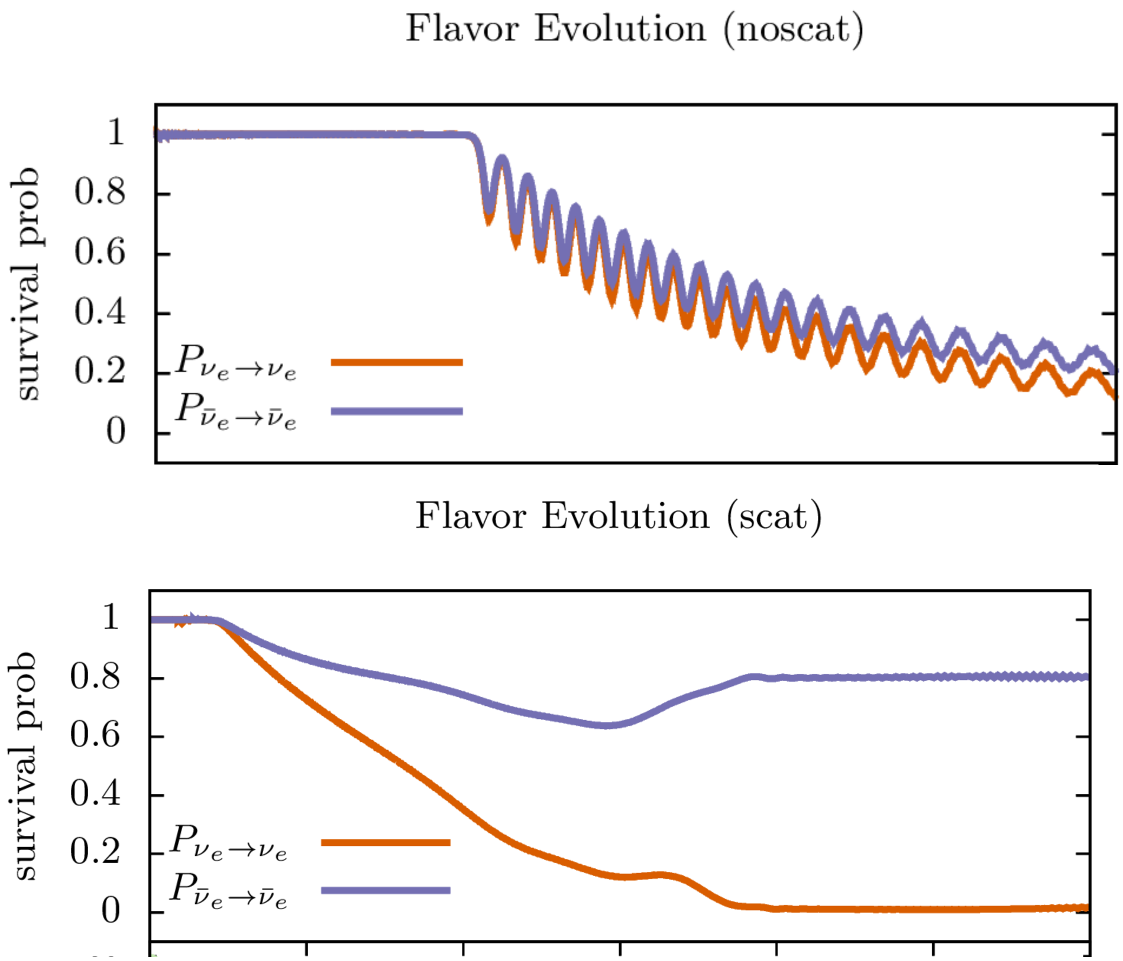
Gava et al 2009

Collective oscillations for supernova nucleosynthesis



Duan et al 2011

Transformation is sensitive to conditions, approximations



Collective

Matter-neutrino Resonance

Deaton et al 2018

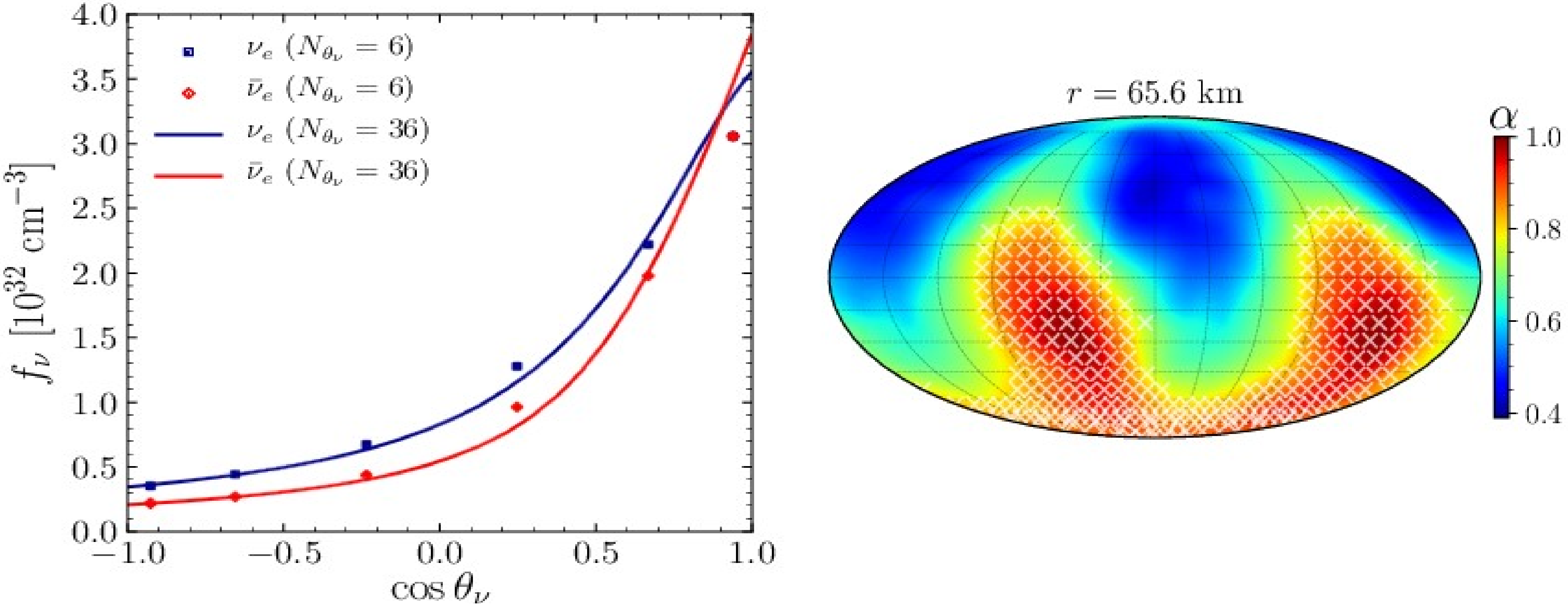
A few of the open issues in supernova neutrino oscillations

“Fast flavor” oscillations may occur close to the decoupling surface – how to follow these very small scale oscillations consistently

Energy, direction, and flavor changing collisions need to be included self consistently

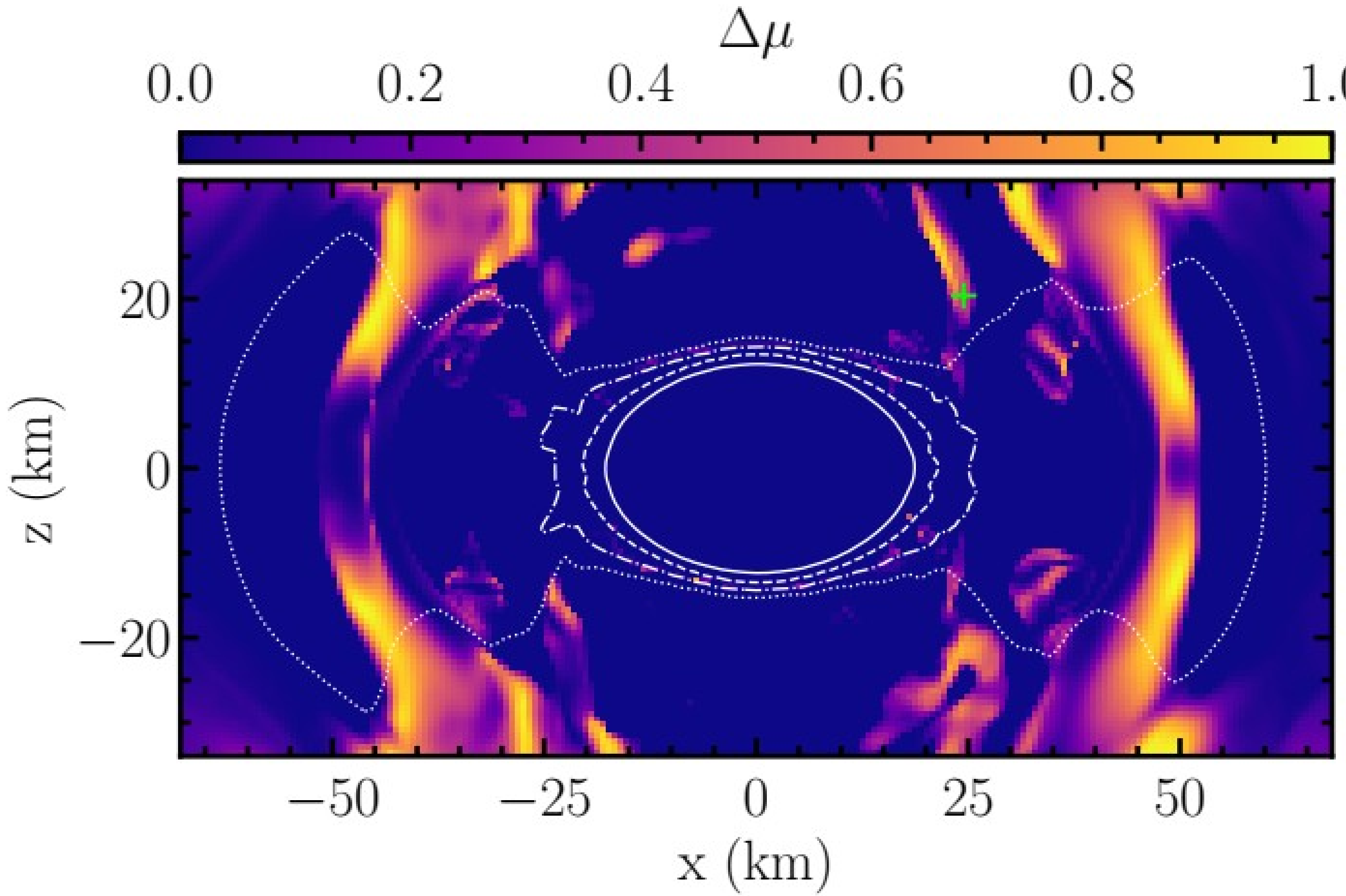
We use mean field – does that make a difference?

Fast oscillation instability correlated with “crossings”



Figures from Abbar et al 2018,
supernova model from Sumiyoshi et al

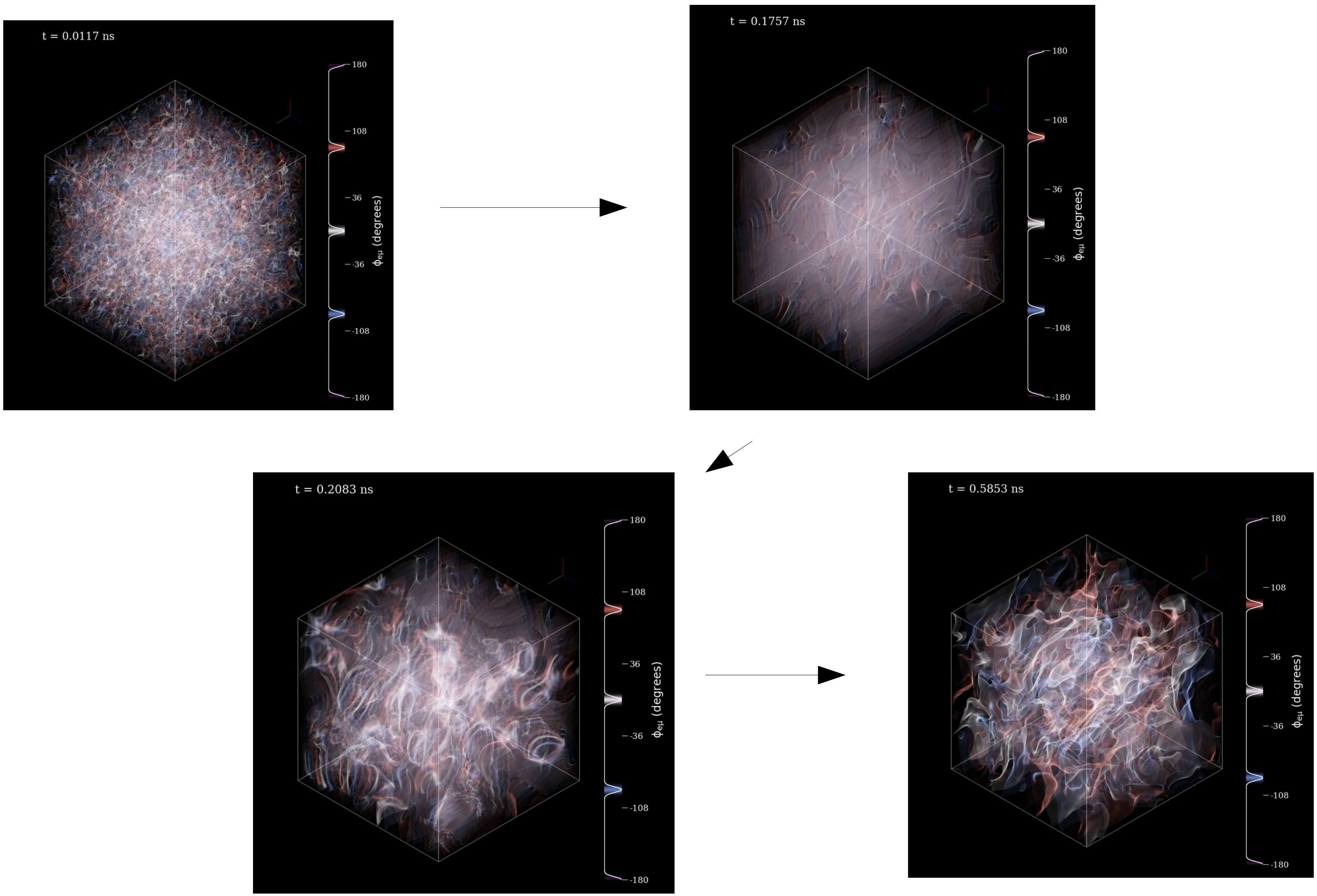
Crossings in BNS remnant



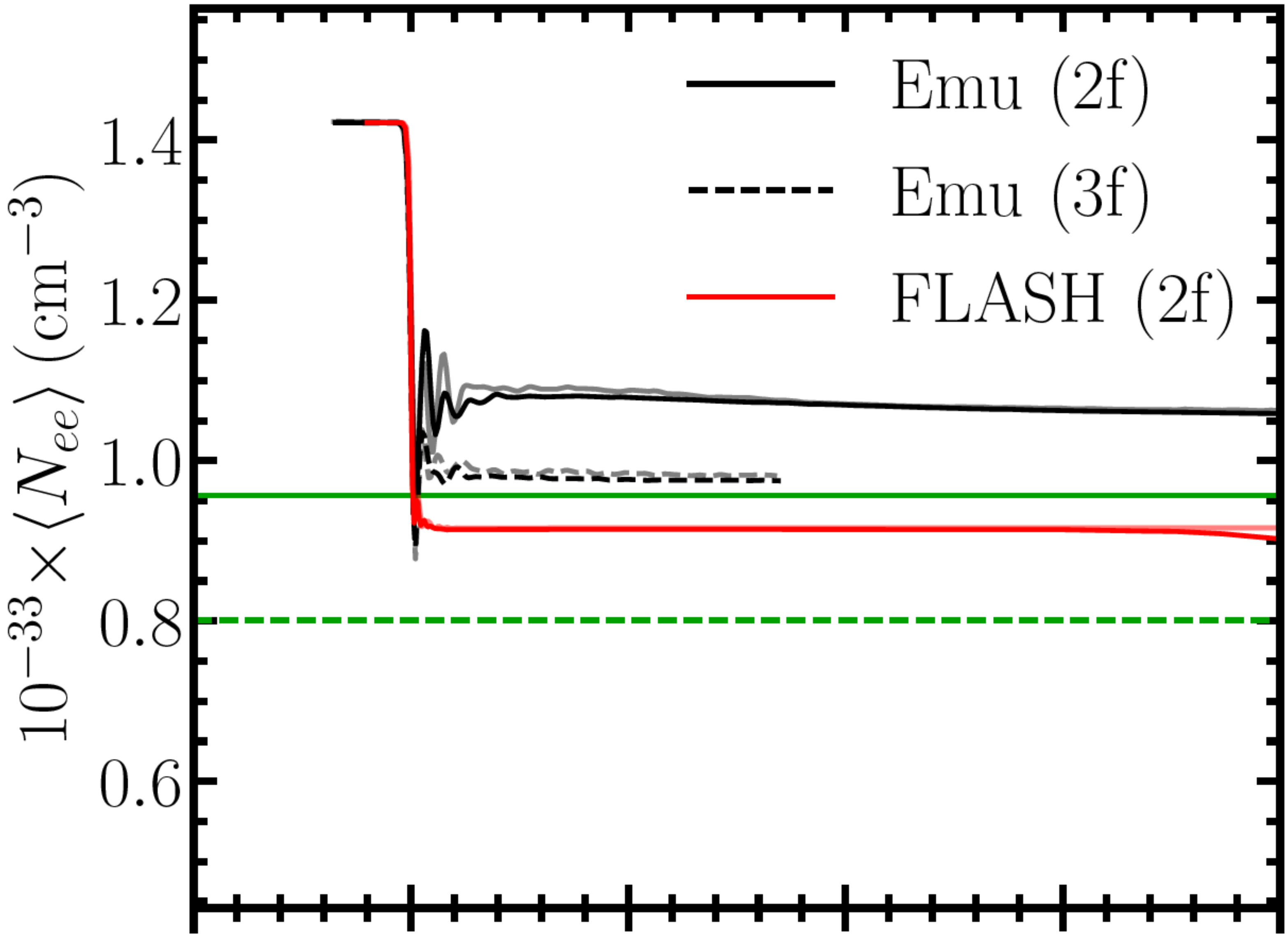
Grohs, Richers et al 2022

Fast flavor oscillations above a BNS merger with moments using FLASH

(Grohs et al 2022)



Growth and saturation, BNS, moments vs PIC



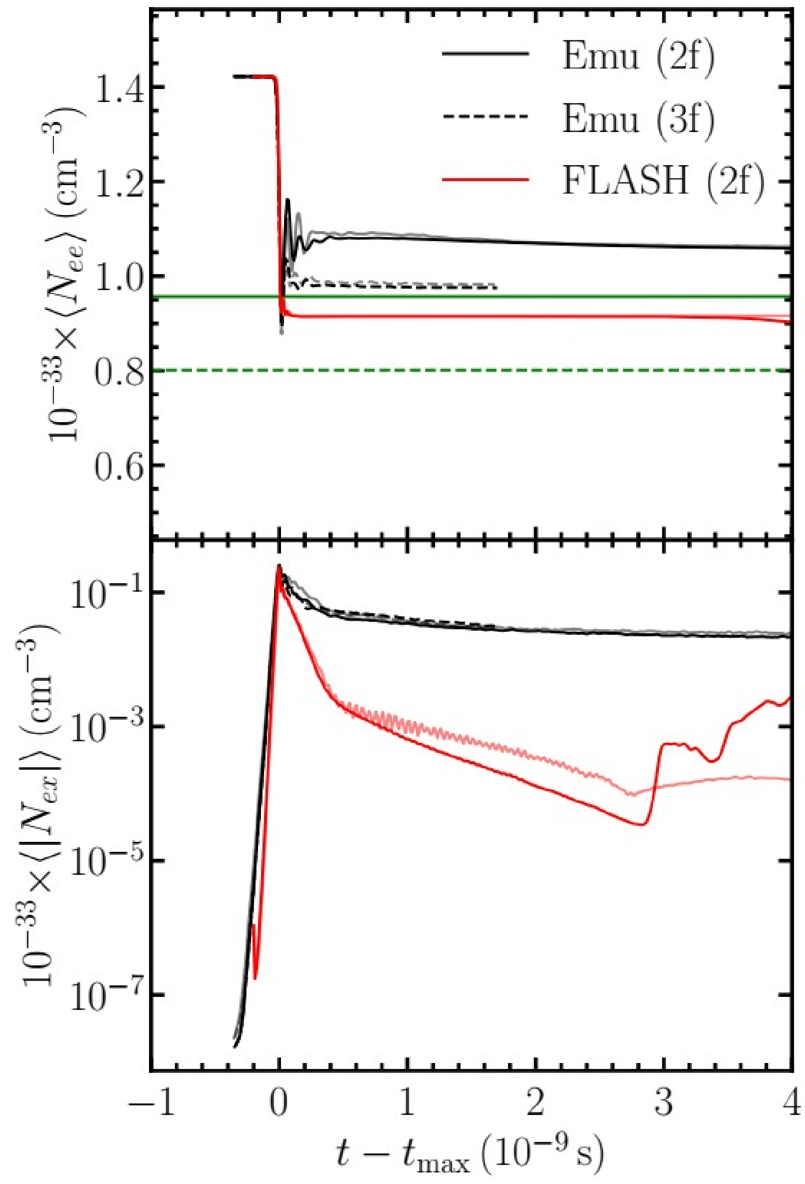
1500 neutrinos per cell

Flavor equilibration

2 moments per cell

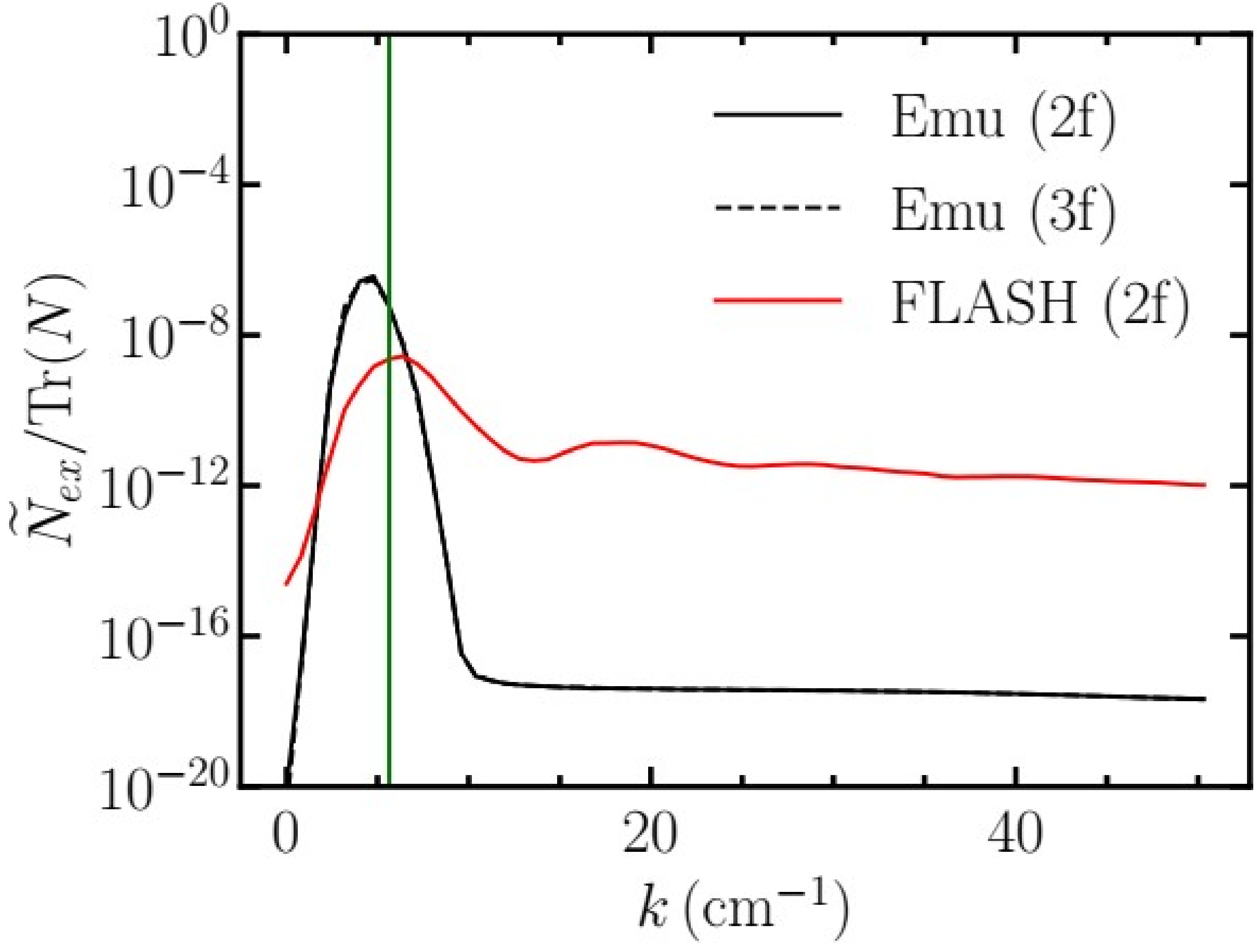
Grohs et al 2022

Growth and saturation, BNS, moments vs PIC



Grohs et al 2022

Fourier transform
BNS, moments vs PIC



Collisions

$$i\frac{D\rho}{Dt} = [\mathbf{H}, \rho] + iC$$

$$i\frac{D\bar{\rho}}{Dt} = [\bar{\mathbf{H}}, \bar{\rho}] + i\bar{C}$$

Collision term

Convective derivative

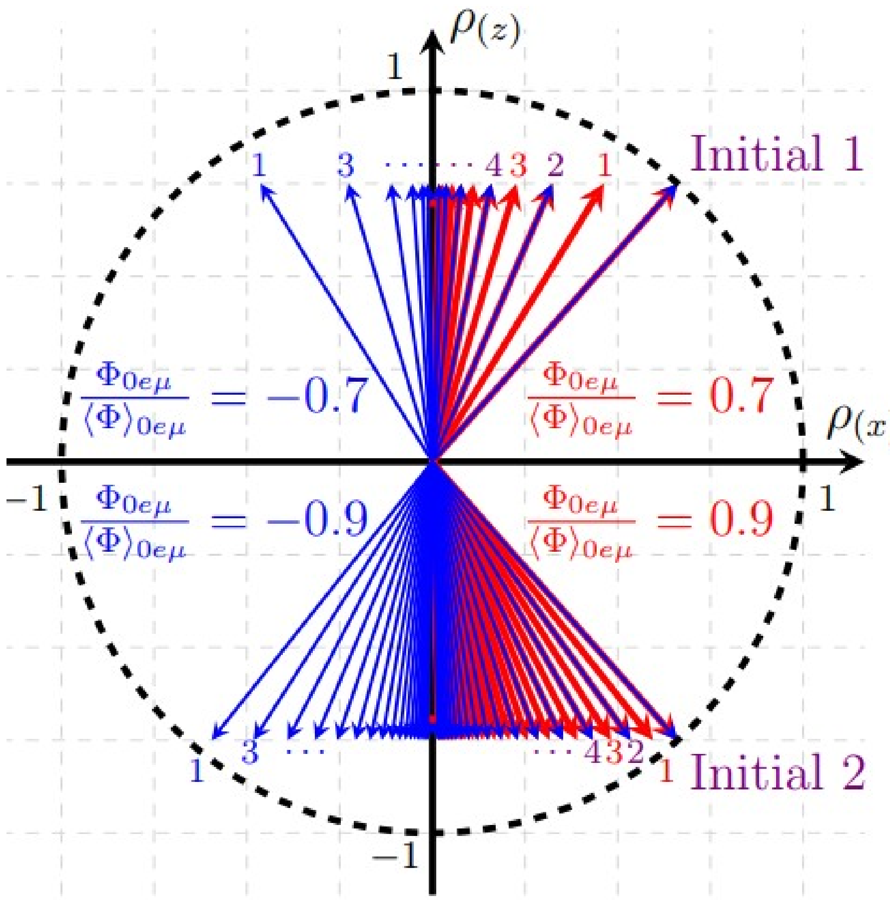
Hamiltonian

Collisions

Collisions: scatterings which change energy, momentum, type of particle

Collisions damp out “mixed” states and send the neutrino system toward pure flavor states (or not! Shalgar et al, Johns et al)

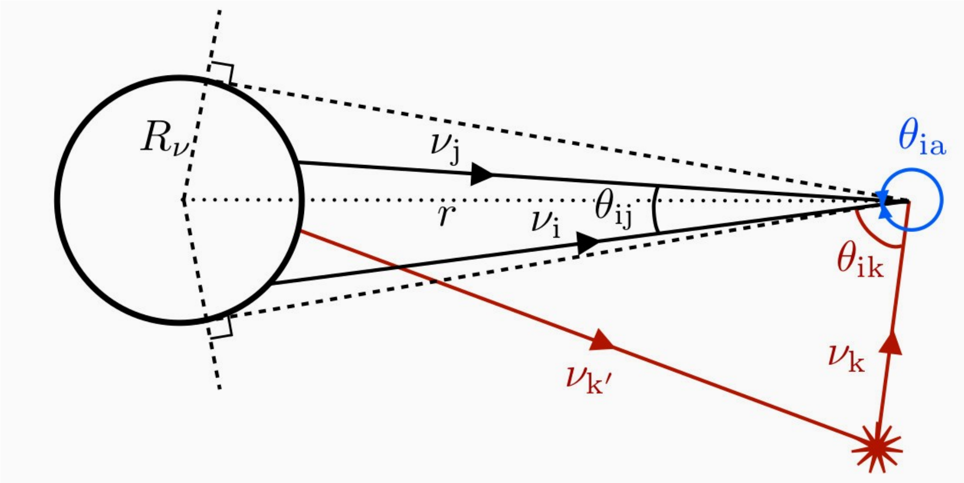
A neutrino in a mixed state under the influence of collisions

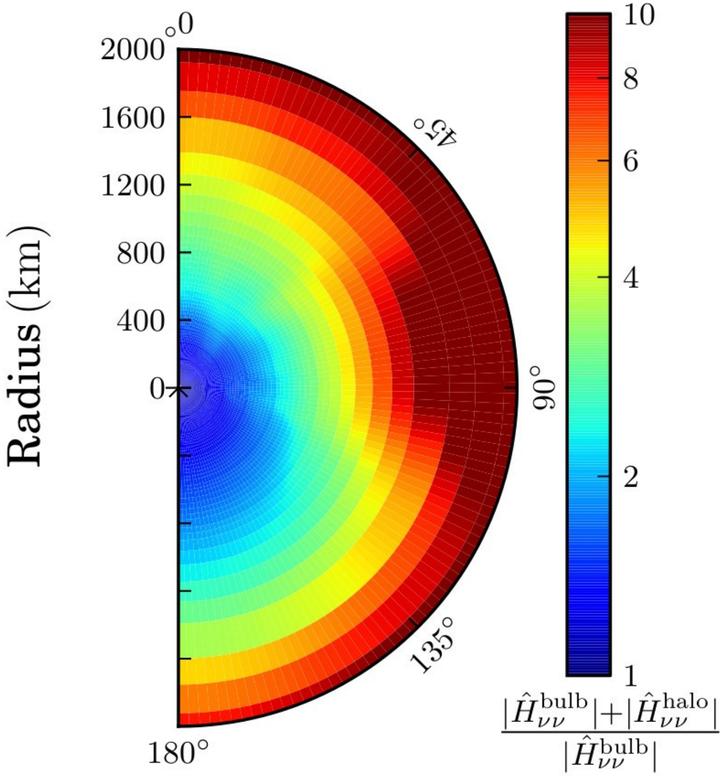


Evolution of flavor vector due to collisions, Fig. from Richers et al, 2019

Halo effect, collisions matter

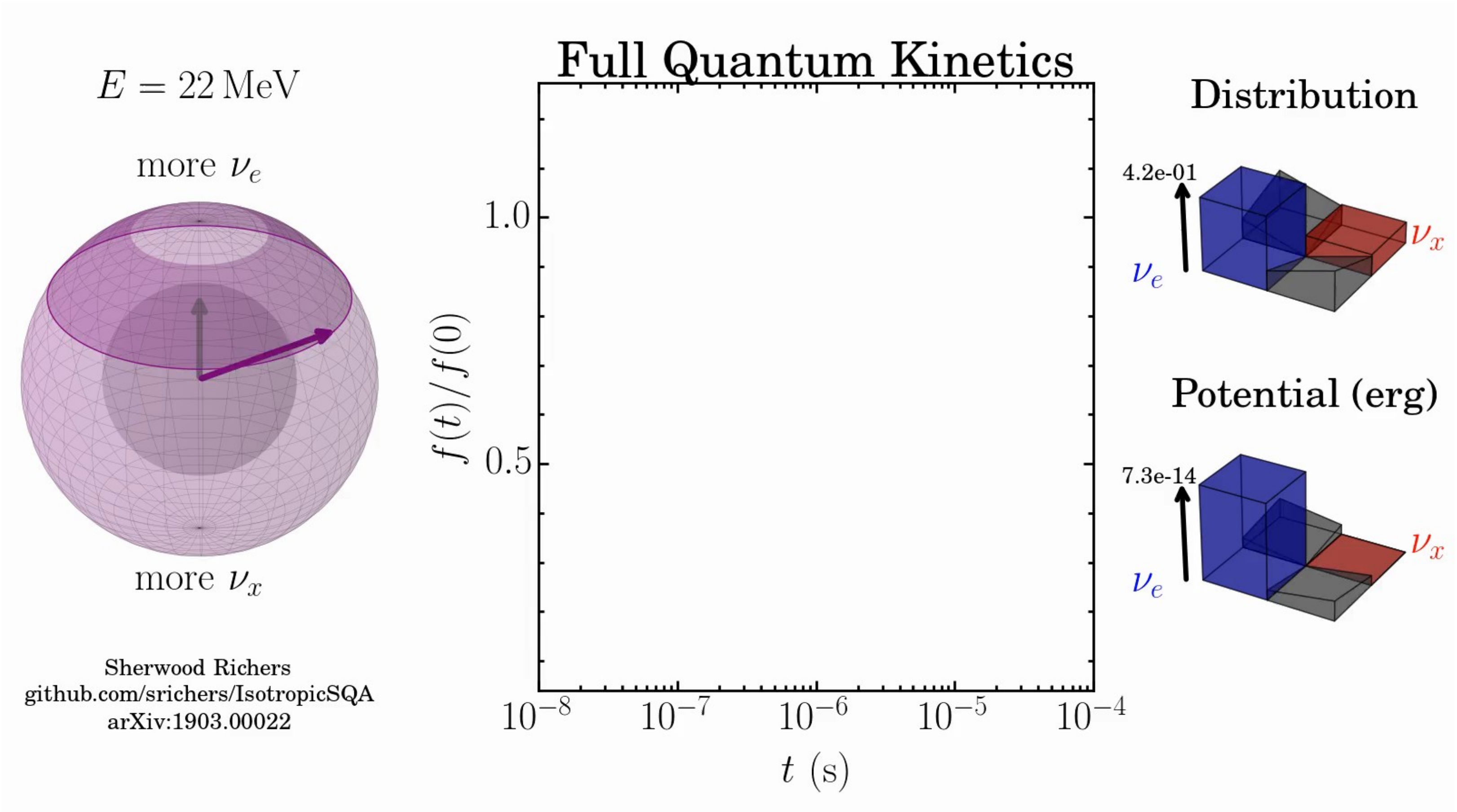
Significant numbers of neutrinos can scatter "backward"



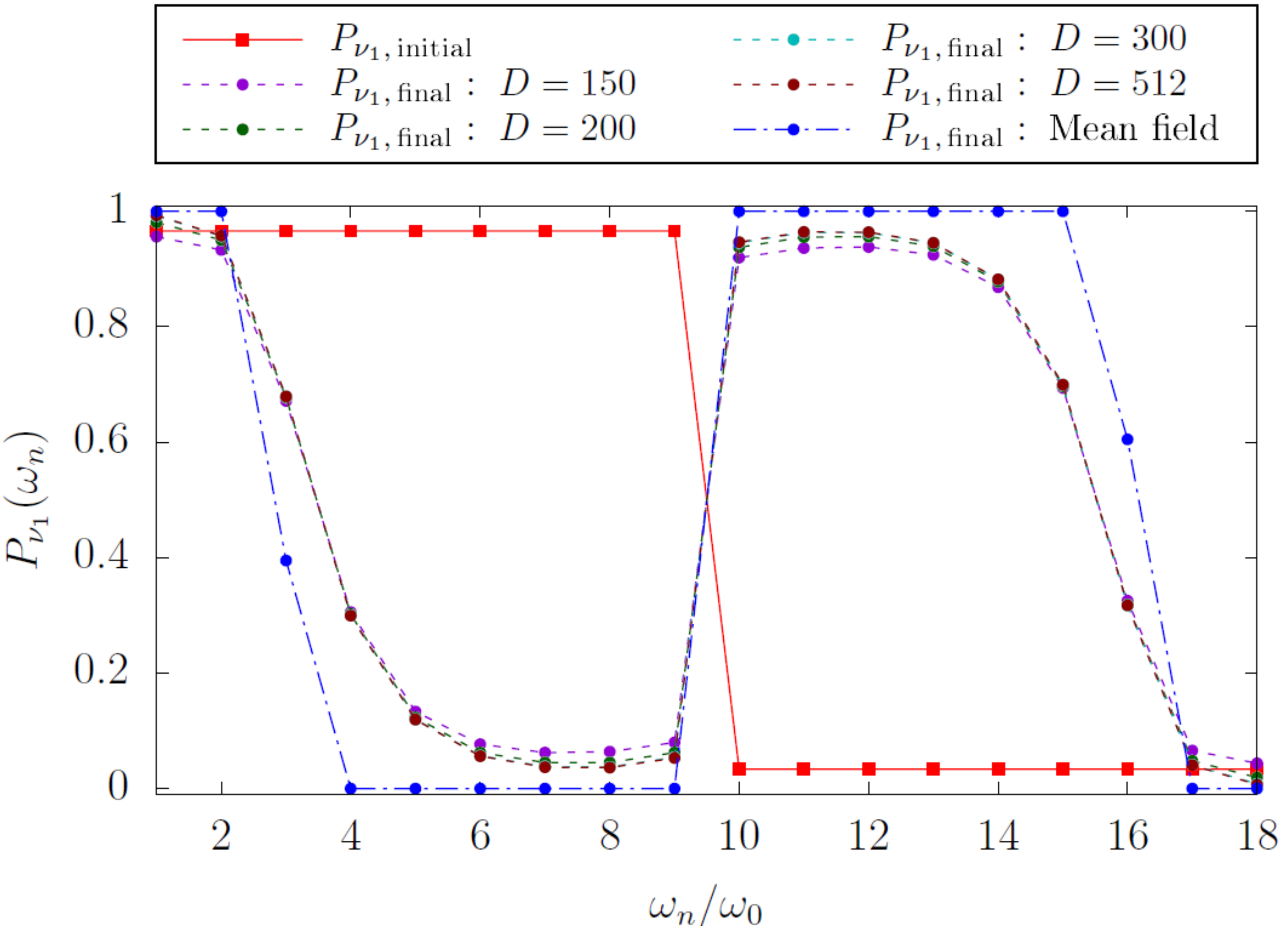


Cherry et al 2012

Oscillations with collisions, isotropic



Beyond mean field



Cervia, Siwach et al 2022

Conclusions

We need to understand neutrinos in astrophysical systems to accurately predict observables including element synthesis, neutrino signals

Involves solving the quantum kinetic equations in astrophysical environments

Starting to make progress on this by understanding fast flavor, making efforts to include full QKEs, understand the usefulness of the mean field approximation

To keep mind: Astrophysical objects will make better laboratories for neutrino physics if we make progress on understanding systems with large numbers of neutrinos